



# Improved Measurement of Muon Antineutrino Disappearance in MINOS

## Citation

Adamson, P., D. S. Ayres, C. Backhouse, G. Barr, M. Bishai, A. Blake, G. J. Bock, et al. 2012. Improved Measurement of Muon Antineutrino Disappearance in MINOS. Physical Review Letters 108, no. 19: 191801.

## Published Version

10.1103/PhysRevLett.108.191801

## Permanent link

<http://nrs.harvard.edu/urn-3:HUL.InstRepos:11870348>

## Terms of Use

This article was downloaded from Harvard University's DASH repository, and is made available under the terms and conditions applicable to Other Posted Material, as set forth at <http://nrs.harvard.edu/urn-3:HUL.InstRepos:dash.current.terms-of-use#LAA>

## Share Your Story

The Harvard community has made this article openly available.  
Please share how this access benefits you. [Submit a story](#).

[Accessibility](#)

## Improved Measurement of Muon Antineutrino Disappearance in MINOS

P. Adamson,<sup>7</sup> D. S. Ayres,<sup>1</sup> C. Backhouse,<sup>18</sup> G. Barr,<sup>18</sup> M. Bishai,<sup>3</sup> A. Blake,<sup>5</sup> G. J. Bock,<sup>7</sup> D. J. Boehnlein,<sup>7</sup> D. Bogert,<sup>7</sup> S. V. Cao,<sup>26</sup> S. Childress,<sup>7</sup> J. A. B. Coelho,<sup>6</sup> L. Corwin,<sup>12</sup> D. Cronin-Hennessy,<sup>15</sup> I. Z. Danko,<sup>19</sup> J. K. de Jong,<sup>18</sup> N. E. Devenish,<sup>24</sup> M. V. Diwan,<sup>3</sup> C. O. Escobar,<sup>6</sup> J. J. Evans,<sup>14</sup> E. Falk,<sup>24</sup> G. J. Feldman,<sup>9</sup> M. V. Frohne,<sup>10</sup> H. R. Gallagher,<sup>27</sup> R. A. Gomes,<sup>8</sup> M. C. Goodman,<sup>1</sup> P. Gouffon,<sup>21</sup> N. Graf,<sup>11</sup> R. Gran,<sup>16</sup> K. Grzelak,<sup>28</sup> A. Habig,<sup>16</sup> J. Hartnell,<sup>24</sup> R. Hatcher,<sup>7</sup> A. Himmel,<sup>4</sup> A. Holin,<sup>14</sup> X. Huang,<sup>1</sup> J. Hylen,<sup>7</sup> G. M. Irwin,<sup>23</sup> Z. Isvan,<sup>19</sup> D. E. Jaffe,<sup>3</sup> C. James,<sup>7</sup> D. Jensen,<sup>7</sup> T. Kafka,<sup>27</sup> S. M. S. Kasahara,<sup>15</sup> G. Koizumi,<sup>7</sup> S. Kopp,<sup>26</sup> M. Kordosky,<sup>29</sup> A. Kreymer,<sup>7</sup> K. Lang,<sup>26</sup> J. Ling,<sup>3,22</sup> P. J. Litchfield,<sup>15,20</sup> L. Loiacono,<sup>26</sup> P. Lucas,<sup>7</sup> W. A. Mann,<sup>27</sup> M. L. Marshak,<sup>15</sup> M. Mathis,<sup>29</sup> N. Mayer,<sup>12</sup> R. Mehdiyev,<sup>26</sup> J. R. Meier,<sup>15</sup> M. D. Messier,<sup>12</sup> D. G. Michael,<sup>4,\*</sup> W. H. Miller,<sup>15</sup> S. R. Mishra,<sup>22</sup> J. Mitchell,<sup>5</sup> C. D. Moore,<sup>7</sup> L. Muallem,<sup>4</sup> S. Mufson,<sup>12</sup> J. Musser,<sup>12</sup> D. Naples,<sup>19</sup> J. K. Nelson,<sup>29</sup> H. B. Newman,<sup>4</sup> R. J. Nichol,<sup>14</sup> J. A. Nowak,<sup>15</sup> W. P. Oliver,<sup>27</sup> M. Orchanian,<sup>4</sup> R. B. Pahlka,<sup>7</sup> J. Paley,<sup>1,12</sup> R. B. Patterson,<sup>4</sup> G. Pawloski,<sup>15,23</sup> S. Phan-Budd,<sup>1</sup> R. K. Plunkett,<sup>7</sup> X. Qiu,<sup>23</sup> A. Radovic,<sup>14</sup> J. Ratchford,<sup>26</sup> B. Rebel,<sup>7</sup> C. Rosenfeld,<sup>22</sup> H. A. Rubin,<sup>11</sup> M. C. Sanchez,<sup>13,19</sup> J. Schneps,<sup>27</sup> A. Schreckenberger,<sup>15</sup> P. Schreiner,<sup>1</sup> R. Sharma,<sup>7</sup> A. Sousa,<sup>9</sup> M. Strait,<sup>15</sup> N. Tagg,<sup>17</sup> R. L. Talaga,<sup>1</sup> J. Thomas,<sup>14</sup> M. A. Thomson,<sup>5</sup> G. Tinti,<sup>18</sup> R. Toner,<sup>5</sup> D. Torretta,<sup>7</sup> G. Tzanakos,<sup>2</sup> J. Urheim,<sup>12</sup> P. Vahle,<sup>29</sup> B. Viren,<sup>3</sup> J. J. Walding,<sup>29</sup> A. Weber,<sup>18,20</sup> R. C. Webb,<sup>25</sup> C. White,<sup>11</sup> L. Whitehead,<sup>3</sup> S. G. Wojcicki,<sup>23</sup> and R. Zwaska<sup>7</sup>

(The MINOS Collaboration)

<sup>1</sup>Argonne National Laboratory, Argonne, Illinois 60439, USA<sup>2</sup>Department of Physics, University of Athens, GR-15771 Athens, Greece<sup>3</sup>Brookhaven National Laboratory, Upton, New York 11973, USA<sup>4</sup>Lauritsen Laboratory, California Institute of Technology, Pasadena, California 91125, USA<sup>5</sup>Cavendish Laboratory, University of Cambridge, Madingley Road, Cambridge CB3 0HE, United Kingdom<sup>6</sup>Universidade Estadual de Campinas, IFGW-UNICAMP, CP 6165, 13083-970, Campinas, SP, Brazil<sup>7</sup>Fermi National Accelerator Laboratory, Batavia, Illinois 60510, USA<sup>8</sup>Instituto de Física, Universidade Federal de Goiás, CP 131, 74001-970, Goiânia, GO, Brazil<sup>9</sup>Department of Physics, Harvard University, Cambridge, Massachusetts 02138, USA<sup>10</sup>Holy Cross College, Notre Dame, Indiana 46556, USA<sup>11</sup>Department of Physics, Illinois Institute of Technology, Chicago, Illinois 60616, USA<sup>12</sup>Indiana University, Bloomington, Indiana 47405, USA<sup>13</sup>Department of Physics and Astronomy, Iowa State University, Ames, Iowa 50011 USA<sup>14</sup>Department of Physics and Astronomy, University College London, Gower Street, London WC1E 6BT, United Kingdom<sup>15</sup>University of Minnesota, Minneapolis, Minnesota 55455, USA<sup>16</sup>Department of Physics, University of Minnesota-Duluth, Duluth, Minnesota 55812, USA<sup>17</sup>Otterbein College, Westerville, Ohio 43081, USA<sup>18</sup>Subdepartment of Particle Physics, University of Oxford, Oxford OX1 3RH, United Kingdom<sup>19</sup>Department of Physics and Astronomy, University of Pittsburgh, Pittsburgh, Pennsylvania 15260, USA<sup>20</sup>Rutherford Appleton Laboratory, Science and Technologies Facilities Council, OX11 0QX, United Kingdom<sup>21</sup>Instituto de Física, Universidade de São Paulo, CP 66318, 05315-970, São Paulo, SP, Brazil<sup>22</sup>Department of Physics and Astronomy, University of South Carolina, Columbia, South Carolina 29208, USA<sup>23</sup>Department of Physics, Stanford University, Stanford, California 94305, USA<sup>24</sup>Department of Physics and Astronomy, University of Sussex, Falmer, Brighton BN1 9QH, United Kingdom<sup>25</sup>Physics Department, Texas A&M University, College Station, Texas 77843, USA<sup>26</sup>Department of Physics, University of Texas at Austin, 1 University Station C1600, Austin, Texas 78712, USA<sup>27</sup>Physics Department, Tufts University, Medford, Massachusetts 02155, USA<sup>28</sup>Department of Physics, University of Warsaw, Hoża 69, PL-00-681 Warsaw, Poland<sup>29</sup>Department of Physics, College of William & Mary, Williamsburg, Virginia 23187, USA

(Received 13 February 2012; published 11 May 2012)

We report an improved measurement of  $\bar{\nu}_\mu$  disappearance over a distance of 735 km using the MINOS detectors and the Fermilab Main Injector neutrino beam in a  $\bar{\nu}_\mu$ -enhanced configuration. From a total exposure of  $2.95 \times 10^{20}$  protons on target, of which 42% have not been previously analyzed, we make the most precise measurement of  $\Delta\bar{m}^2 = [2.62_{-0.28}^{+0.31}(\text{stat}) \pm 0.09(\text{syst})] \times 10^{-3} \text{ eV}^2$  and constrain the  $\bar{\nu}_\mu$  mixing angle  $\sin^2(2\bar{\theta}) > 0.75$  (90% C.L.). These values are in agreement with  $\Delta m^2$  and  $\sin^2(2\theta)$  measured

for  $\nu_\mu$ , removing the tension reported in [P. Adamson *et al.* (MINOS), *Phys. Rev. Lett.* **107**, 021801 (2011)].

DOI: 10.1103/PhysRevLett.108.191801

PACS numbers: 14.60.Pq, 14.60.Lm, 29.27.-a

Observations of neutrinos and antineutrinos created in the Sun, the Earth's atmosphere, reactors, and accelerators provide strong evidence [1–9] that neutrinos undergo transitions between their flavor eigenstates ( $\nu_e$ ,  $\nu_\mu$ ,  $\nu_\tau$ ) as they propagate. These transitions can occur due to quantum mechanical mixing between the neutrino flavor and mass ( $\nu_1$ ,  $\nu_2$ ,  $\nu_3$ ) eigenstates. The mixing may be parametrized with a unitary matrix  $U_{\text{PMNS}}$  [10] which is typically expressed in terms of three mixing angles  $\theta_{12}$ ,  $\theta_{23}$ ,  $\theta_{13}$  and a charge-parity ( $CP$ ) violating phase  $\delta$ . This interpretation, referred to as “neutrino oscillations,” requires that neutrinos have mass and motivates extensions to the standard model (SM) of particle physics. Extensions that explain the origin of neutrino masses, for example, the addition of right-handed sterile neutrinos  $\nu_R$  [11], may also explain the baryon asymmetry [12] of the Universe.

The  $CPT$  symmetry of the SM requires that  $\nu_\mu$  and  $\bar{\nu}_\mu$  have the same masses and mixing parameters. In vacuum, the probability  $P(\nu_\mu \rightarrow \nu_\mu)$  that a  $\nu_\mu$  is detected after a distance  $L$  as a  $\nu_\mu$  (rather than a  $\nu_e$  or  $\nu_\tau$ ) must be equal to the corresponding probability  $P(\bar{\nu}_\mu \rightarrow \bar{\nu}_\mu)$  for antineutrinos. For a  $\nu_\mu$  with energy  $E$ , the probability may be written as

$$P(\nu_\mu \rightarrow \nu_\mu) = 1 - \sin^2(2\theta) \sin^2\left(\frac{1.267 \Delta m^2 [\text{eV}^2] L [\text{km}]}{E [\text{GeV}]}\right), \quad (1)$$

where  $\Delta m^2$  and  $\sin^2(2\theta)$  are effective parameters that are functions of the angles parametrizing  $U_{\text{PMNS}}$  and the differences in the squared masses  $\Delta m_{ij}^2 = m_i^2 - m_j^2$  of the  $\nu_1$ ,  $\nu_2$ , and  $\nu_3$  states. Experiments have demonstrated  $|\Delta m_{31}^2| \gg |\Delta m_{21}^2|$  [13]. In the limiting case that  $\theta_{13} \approx 0$  we have  $\theta \approx \theta_{23}$  and  $|\Delta m^2| \approx \sin^2(2\theta_{12})|\Delta m_{31}^2| + \cos^2(2\theta_{12})|\Delta m_{32}^2|$  [14,15]. Muon antineutrino oscillations are described by an equation which has the same form as Eq. (1) with parameters  $\Delta \bar{m}^2$  and  $\sin^2(2\bar{\theta})$ . The extended SM predicts  $\Delta \bar{m}^2 = \Delta m^2$  and  $\sin^2(2\bar{\theta}) = \sin^2(2\theta)$  for vacuum oscillations [16,17]. Observation of  $P(\nu_\mu \rightarrow \nu_\mu) \neq P(\bar{\nu}_\mu \rightarrow \bar{\nu}_\mu)$  would therefore be evidence for physics beyond the SM, such as neutrino interactions in the Earth's crust that do not conserve lepton flavor.

In this Letter, we describe a measurement of  $P(\bar{\nu}_\mu \rightarrow \bar{\nu}_\mu)$  conducted over a baseline  $L = 735$  km using a  $\bar{\nu}_\mu$ -enhanced beam with a peak energy of 3 GeV. The beam was produced by directing 120 GeV/c protons from the Fermilab Main Injector onto a graphite target to produce  $\pi/K$  mesons that decay to produce neutrinos. Two

magnetic horns focus the mesons, allowing us to control the energy spectrum and  $\nu/\bar{\nu}$  content of the beam.

The neutrino beam is pointed towards two detectors, referred to as near and far. The 980 ton near detector (ND) measures the  $\nu_\mu$  and  $\bar{\nu}_\mu$  content of the beam as a function of energy at a distance of 1.04 km from the  $\pi/K$  production target. The 5.4 kton far detector (FD) is located in the Soudan Underground Laboratory, 734 km from the ND, and remeasures the beam composition. The neutrino detectors are steel-scintillator, tracking-sampling calorimeters optimized to identify and measure the energy of muon neutrinos and antineutrinos and reject backgrounds from neutral-current (NC) and  $\nu_e$  interactions [18]. The detectors are magnetized with an average field of 1.3 T to distinguish  $\nu_\mu$  from  $\bar{\nu}_\mu$  based on the charge of the  $\mu$  produced in weak interactions.

We previously reported  $\nu_\mu$  oscillations with an energy dependence consistent with Eq. (1) and  $\Delta m^2 = [2.32^{+0.12}_{-0.08}(\text{stat} + \text{syst})] \times 10^{-3} \text{ eV}^2$ ,  $\sin^2(2\theta) > 0.90$  (90% C.L.) [3]. The measurements utilized  $7.25 \times 10^{20}$  protons on target (POT) of data collected between 2005–2009 with a  $\nu_\mu$ -enhanced beam [19]. Measurements made by Super-Kamiokande [4] and T2K [20] are in good agreement with our values.

In 2009–2010 we collected  $1.71 \times 10^{20}$  POT in a  $\bar{\nu}_\mu$ -enhanced beam [19] created by reversing the polarity of the horns. The magnetic fields in the FD and ND were also reversed to focus the  $\mu^+$  created in  $\bar{\nu}_\mu$  interactions. These antineutrino data also exhibited oscillations in agreement with Eq. (1), but with parameters  $\Delta \bar{m}^2 = [3.36^{+0.46}_{-0.40}(\text{stat}) \pm 0.06(\text{syst})] \times 10^{-3} \text{ eV}^2$  and  $\sin^2(2\bar{\theta}) = 0.86 \pm 0.11(\text{stat}) \pm 0.01(\text{syst})$  [21]. We found no systematic effects which could explain the difference between the  $\nu_\mu$  and  $\bar{\nu}_\mu$  parameters. Assuming identical true values for  $(\Delta \bar{m}^2, \sin^2(2\bar{\theta}))$  and  $(\Delta m^2, \sin^2(2\theta))$ , we calculated that such a difference would occur by random chance about 2% of the time. To clarify the situation, we collected an additional  $1.24 \times 10^{20}$  POT with the  $\bar{\nu}_\mu$ -enhanced beam during 2010–2011. We have also updated the analysis to improve the sensitivity to  $[\Delta \bar{m}^2, \sin^2(2\bar{\theta})]$ , reduce uncertainties due to Monte Carlo (MC) modeling in the ND, and increase the similarity to the  $\nu_\mu$  oscillation analysis.

We isolate a sample of  $\nu_\mu$  and  $\bar{\nu}_\mu$  charged-current (CC)  $\nu_\mu N \rightarrow \mu X$  events by searching for interaction vertices inside our detectors with a muon track and possible hadronic activity from the recoil system  $X$ . We reject hadron tracks reconstructed in NC events by combining four topological variables describing track properties into a single discriminant variable,  $\mu\text{ID}$ , using a  $k$ -nearest-neighbor

(kNN) technique [22]. The kNN algorithm calculates the distance in the four-variable space between each measured event and an ensemble of simulated events; the output is the fraction of signal in the  $k = 80$  closest MC events. This discriminant was used in our previous analyses [3,21] and, as shown in Fig. 1(a), is well modeled by our MC simulation. We maximize the statistical sensitivity to  $\Delta\bar{m}^2$  by requiring  $\mu\text{ID} > 0.3$ , which results in a MC estimated efficiency (purity) of 90.7% (99.0%) at the ND and 91.6% (99.0%) at the FD. We then discriminate  $\nu_\mu$  from  $\bar{\nu}_\mu$  on an event-by-event basis by analyzing the track curvature in the detector's magnetic field. Figure 1(b) shows the track charge/momentum ( $q/p$ ) divided by its uncertainty [ $\sigma(q/p)$ ], as determined by our track reconstruction algorithm. We select  $\bar{\nu}_\mu$  and reject  $\nu_\mu$  by requiring  $\frac{q/p}{\sigma(q/p)} > 0$  with a MC estimated efficiency (purity) of 98.4% (94.7%) in the ND and 98.8% (95.1%) in the FD. The  $\nu_\mu$  background accepted by the selection is predominantly due to high energy muons with small curvature.

We reconstruct the neutrino energy by summing muon and hadronic shower energies. The muon energy is measured using track range and curvature. We reconstruct the hadronic shower energy using three variables: the sum of the reconstructed energy deposited by showers that start within 1 m of the track vertex, the sum of the energy in the two largest showers reconstructed in the event, and the length of the longest shower. We use these three variables

in a second kNN algorithm and estimate the shower energy as the mean true hadronic energy of the  $k = 400$  closest MC events. This technique improves the hadronic energy resolution when compared to a method which uses only the energy deposited by the largest shower, increases the statistical sensitivity to  $\Delta\bar{m}^2$  by 10%, and was previously used to analyze the  $\nu_\mu$  disappearance [3,23].

Data from the ND are used to predict the neutrino energy distribution at the FD. Though both detectors have the same segmentation and very similar average magnetic fields, for economic reasons the ND is smaller and asymmetric about the magnetic field coil and is more coarsely instrumented with scintillator in the downstream “muon spectrometer” region [18]. In addition, the ND coil occupies a larger fractional area than the FD coil and more muons enter it. In the ND data, we observe a reconstruction failure rate of 6.1%, mostly associated with tracks entering the coil region, but the MC simulation predicts 4.2%. Previously, we dealt with this issue by assigning a systematic error. Now, we remove ND events with a track that ends less than 60 cm from the coil. We also remove events with a track that ends on the side of the coil opposite the beam centroid. The new event selection decreases the efficiency to 53% in the ND, but reduces the data and MC failure rates to 1.4% and 0.9%, respectively. The selected sample contains the same classes of neutrino scattering processes as are present at the FD, and our results are not significantly more vulnerable to cross-section uncertainties. We applied the new selection and shower energy reconstruction to the 2009–2010 data and found that the best fit parameters shifted by only  $\delta(\Delta\bar{m}^2) = [+1.0 \times 10^{-4}] \text{eV}^2$  and  $\delta[\sin^2(2\bar{\theta})] = -3.6 \times 10^{-2}$ .

A total of  $2.98 \times 10^{20}$  protons were delivered to the graphite target during the data-taking periods. We impose a series of data and beam quality [2,24,25] requirements which reduce the analyzable exposure to  $2.95 \times 10^{20}$  POT (99.0% live time) at the FD, and  $2.73 \times 10^{20}$  POT (91.8% live time) at the ND. The uncertainty in the live time is negligible, and the high and largely overlapping live time in both detectors assures that our results are not sensitive to beam effects that would cause the number of events per POT to vary.

Under normal conditions, the ND measures about 2400  $\bar{\nu}_\mu$ -CC events per day in the oscillation energy region ( $E_\nu < 6 \text{ GeV}$ ). These data are essential for monitoring the neutrino beam and the quality of the experiment. Figure 2 shows the reconstructed  $\bar{\nu}_\mu$  energy distribution measured in each month during the data-taking periods. With the exception of Feb. 2011 (the last point in each bin in Fig. 2), all months are in statistical agreement, and we expect a constant counting rate per POT at the FD. Part of the February dataset was taken after the neutrino target's cooling system had failed, leaking water into the target canister. This resulted in a decrease in the neutrino flux of 4% from 0–6 GeV when integrated over the entire month.

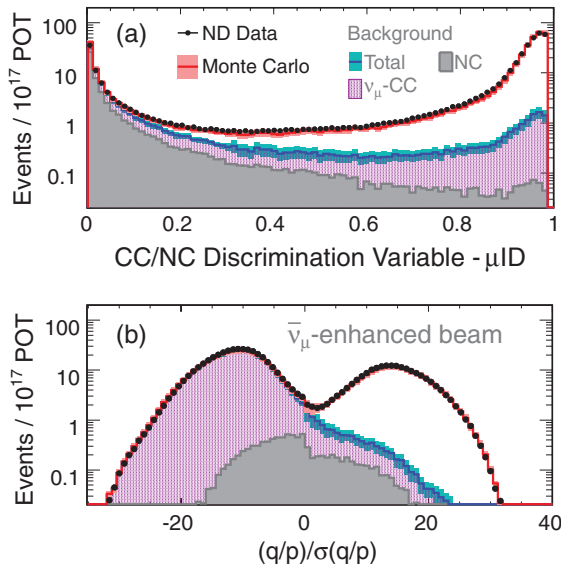


FIG. 1 (color). Event selection variables from the ND: (a) The variable,  $\mu\text{ID}$ , used to select  $\nu_\mu/\bar{\nu}_\mu$ -CC events and reject NC events. (b) The reconstructed track charge/momentum  $q/p$  divided by the uncertainty  $\sigma(q/p)$  reported by the track reconstruction algorithm. In both panels, we have applied all selection criteria except the one on the quantity being shown. Shaded bands show systematic uncertainties. The background histograms are stacked on top of each other.



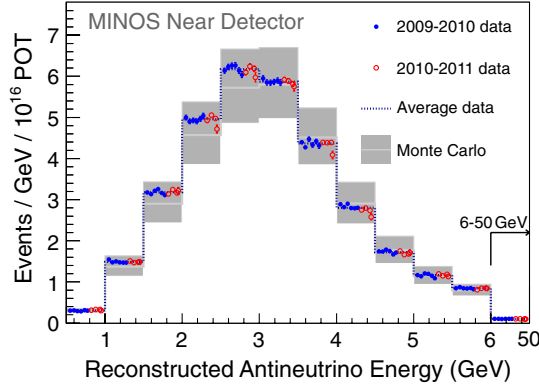


FIG. 2 (color online). The  $\bar{\nu}_\mu$ -CC energy spectrum measured by the ND. The selection procedure described in the text has been applied. Each point represents one month of data. The shaded Monte Carlo band shows the combined effect of systematic uncertainties due to cross sections, flux, energy scale, and other sources.

The decrease is adequately modeled by our beam simulation, and we account for it at the FD using the ND data.

Measurements of the beam position and width at the target, the position of the remnant proton beam at the end of the decay pipe, and the integrated muon flux from meson decays in the pipe, all further indicate that the expected number of  $\bar{\nu}_\mu$ -CC per POT is constant at the FD. Using the proton beam and monitors, we measured the target and horn misalignments as well as the residual magnetic field in the neck of the focusing horns [26]. We conclude that these effects introduce  $\ll 1\%$  uncertainty in the rate at the FD.

Events at the FD are read out in a  $100\ \mu\text{s}$  window surrounding the  $10\ \mu\text{s}$  long beam spill [27]. We select events coincident with the beam spill that have a vertex in the fiducial volume and a track identified as a muon by  $\mu\text{ID}$ . We remove the cosmic-ray background by requiring that the cosine of the angle between the track and beam direction be  $>0.6$ . A total of 521 events satisfy our criteria, 273 in the 2009–2010 data set and 248 in 2010–2011. Using the muon charge, we identify 328 as  $\nu_\mu$  and 193 as  $\bar{\nu}_\mu$ . We apply the selection to  $37.7 \times 10^7$  readout windows taken in anticoincidence with the beam during 2005–2011 and estimate a cosmic-ray background of 0.8 events. We monitor detector and reconstruction performance by measuring the rate of cosmic-ray muons traversing the detector before and after the beam spill. The rates in each data-taking period,  $0.38 \pm 0.03\ \text{Hz}$  in 2009–2010 and  $0.41 \pm 0.04\ \text{Hz}$  in 2010–2011, are consistent.

We predict the  $\bar{\nu}_\mu$  energy spectrum at the FD by first correcting the ND spectrum for inefficiency and backgrounds. We then transfer that spectrum to the FD using a two dimensional “beam matrix” [2]. We predict 273 events if  $\bar{\nu}_\mu$  do not oscillate, including 3.5 NC and 10.8  $\nu_\mu$  CC [28]. The energy spectra are shown in Fig. 3. Oscillations are incorporated into the prediction according to Eq. (1). Maximizing the binned log-likelihood yields

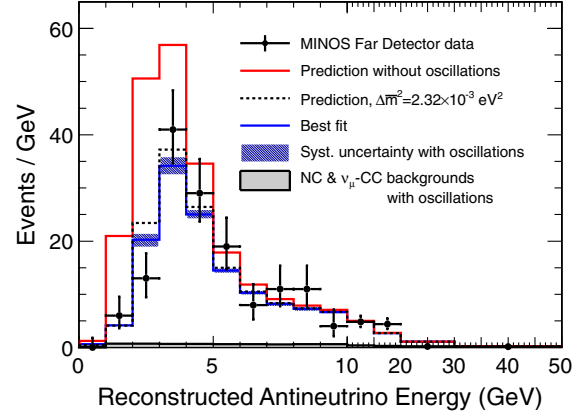


FIG. 3 (color online). The reconstructed energy of FD  $\bar{\nu}_\mu$ -CC events. These data are fit to produce the 2009–2011 contour shown in Fig. 4. The band displays the effect of systematic uncertainties.

$$\begin{aligned}\Delta\bar{m}^2 &= [2.62^{+0.31}_{-0.28}(\text{stat}) \pm 0.09(\text{syst})] \times 10^{-3}\ \text{eV}^2, \\ \sin^2(2\bar{\theta}) &= 0.95^{+0.10}_{-0.11}(\text{stat}) \pm 0.01(\text{syst}), \\ \sin^2(2\bar{\theta}) &> 0.75(90\%\text{C.L.}),\end{aligned}$$

with  $p = 35.3\%$  at best fit [29]. When analyzing the 2010–2011 alone we obtain  $\Delta\bar{m}^2 = [2.26^{+0.27}_{-0.29}(\text{stat}) \pm 0.09(\text{syst})] \times 10^{-3}\ \text{eV}^2$ ,  $\sin^2(2\bar{\theta}) > 0.79$  (90% C.L.) with  $p = 14.5\%$ .

The systematic errors stated above were evaluated by computing the standard deviation in the best fit parameters as the fits were repeated using MC samples that were shifted in accordance with uncertainties in neutrino cross sections, the beam matrix and relative FD normalization, NC and  $\nu_\mu$ -CC backgrounds, the relative FD to ND energy calibration, the absolute muon energy scale, and the absolute hadronic energy scale, including final state hadronic interactions. The input uncertainties are as in our previous analyses [3,21] and were derived from *in situ* data, bench tests of detector and beam components, a test beam experiment, and published neutrino and hadron cross sections [2,30]. We increased the uncertainty in the axial-vector mass from 15% to 30% to account for additional uncertainties in  $\bar{\nu}_\mu$  quasielastic scattering [31].

The  $\bar{\nu}_\mu$  confidence regions shown in Fig. 4 were calculated according to the unified procedure of [32] and incorporate both statistical and systematic uncertainties. Contours for the larger  $\nu_\mu$  data set were produced with two analysis enhancements [3], namely, binning events by energy resolution and treating some systematic effects as nuisance parameters, which do not benefit the lower statistics  $\bar{\nu}_\mu$  analysis presented here. We assess the consistency of the  $\nu_\mu$  and  $\bar{\nu}_\mu$  measurements by doing a joint  $[\Delta\bar{m}^2 = \Delta m^2, \sin^2(2\bar{\theta}) = \sin^2(2\theta)]$  fit to the data to establish the null hypothesis. We then perform four-parameter  $[\Delta\bar{m}^2, \Delta m^2, \sin^2(2\bar{\theta}), \sin^2(2\theta)]$  fits on an ensemble of  $\nu_\mu$

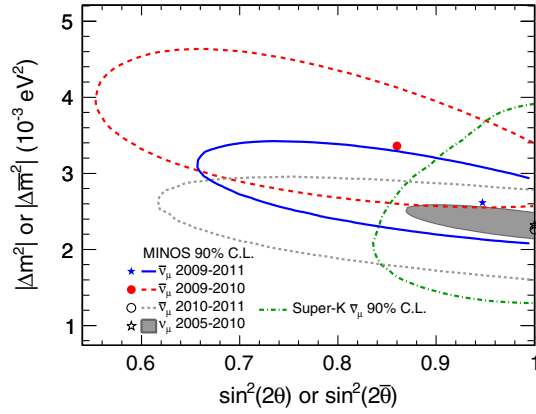


FIG. 4 (color online). Confidence regions calculated by fitting the FD data with an energy spectrum predicted from the ND but modified to incorporate oscillations via Eq. (1). The regions include statistical and systematic uncertainties. The  $\bar{\nu}_\mu$  2009–2011 contour is derived from the energy spectrum shown in Fig. 3.

and  $\bar{\nu}_\mu$  MC experiments generated at the joint best fit. The joint fit had a larger likelihood than  $p = 42\%$  of the four-parameter fits, indicating consistency between  $\nu_\mu$  and  $\bar{\nu}_\mu$ .

In conclusion, we have used a  $\bar{\nu}_\mu$ -enhanced Fermilab accelerator beam and detectors that discriminate  $\nu_\mu$  from  $\bar{\nu}_\mu$  to make the most precise measurement of  $\Delta\bar{m}^2$ . Our results remove the tension reported in [21] and establish consistency between  $\nu_\mu$  and  $\bar{\nu}_\mu$  oscillations at  $L/E \approx 200$  km/GeV.

This work was supported by the U.S. DOE; the United Kingdom STFC; the U.S. NSF; the State and University of Minnesota; the University of Athens, Greece; Brazil's FAPESP, CNPq and CAPES; and the Jeffress Memorial Trust. We are grateful to the Minnesota Department of Natural Resources and the personnel of the Soudan Laboratory and Fermilab.

\*Deceased.

- [1] D. G. Michael *et al.* (MINOS Collaboration), *Phys. Rev. Lett.* **97**, 191801 (2006).
- [2] P. Adamson *et al.* (MINOS Collaboration), *Phys. Rev. D* **77**, 072002 (2008).
- [3] P. Adamson *et al.* (MINOS Collaboration), *Phys. Rev. Lett.* **106**, 181801 (2011).
- [4] Y. Ashie *et al.* (Super-Kamiokande Collaboration), *Phys. Rev. Lett.* **93**, 101801 (2004).
- [5] W. W. M. Allison *et al.* (Soudan-2 Collaboration), *Phys. Rev. D* **72**, 052005 (2005).
- [6] M. Ambrosio *et al.* (MACRO Collaboration), *Eur. Phys. J. C* **36**, 323 (2004).
- [7] M. H. Ahn *et al.* (K2K Collaboration), *Phys. Rev. D* **74**, 072003 (2006).
- [8] Q. R. Ahmad *et al.* (SNO Collaboration), *Phys. Rev. Lett.* **89**, 011301 (2002).
- [9] S. Abe *et al.* (KamLAND Collaboration), *Phys. Rev. Lett.* **100**, 221803 (2008).
- [10] PMNS refers to Pontecorvo, Maki, Nakagawa and Sakata. Z. Maki, M. Nakagawa, and S. Sakata, *Prog. Theor. Phys.* **28**, 870 (1962); B. Pontecorvo, *Zh. Eksp. Teor. Fiz.* **53**, 1717 (1967) [*Sov. Phys. JETP* **26**, 984 (1968)]; V. N. Gribov and B. Pontecorvo, *Phys. Lett. B* **28**, 493 (1969).
- [11] R. N. Mohapatra and G. Senjanovic, *Phys. Rev. Lett.* **44**, 912 (1980).
- [12] M. Fukugita and T. Yanagida, *Phys. Lett. B* **174**, 45 (1986).
- [13] K. Nakamura *et al.* (Particle Data Group), *J. Phys. G* **37**, 075021 (2010).
- [14] S. J. Parke, in *Neutrino Oscillations: Present Status and Future Plans*, edited by J. A. Thomas and P. L. Vahle (World Scientific, Singapore, 2008).
- [15] T. Araki *et al.* (KamLAND Collaboration), *Phys. Rev. Lett.* **94**, 081801 (2005).
- [16] In matter,  $P(\nu_\mu \rightarrow \nu_\mu)$  and  $P(\bar{\nu}_\mu \rightarrow \bar{\nu}_\mu)$  can differ by as much as 0.1% due to  $\nu_\mu \leftrightarrow \nu_e$  mixing and  $\nu_e$  and  $\bar{\nu}_e$  scattering on electrons. This effect is too small to be observed in our experiment.
- [17] L. Wolfenstein, *Phys. Rev. D* **17**, 2369 (1978).
- [18] D. G. Michael *et al.* (MINOS), *Nucl. Instrum. Methods Phys. Res., Sect. A* **596**, 190 (2008).
- [19] Accounting for flux and cross sections, the  $\nu_\mu(\bar{\nu}_\mu)$  enhanced beam consists of 91.7%(58.1%)  $\nu_\mu$ , 7.0%(39.9%)  $\bar{\nu}_\mu$ , and 1.3%(2.0%)  $\nu_e + \bar{\nu}_e$ .
- [20] K. Abe *et al.* (T2K Collaboration), *Phys. Rev. D* **85**, 031103 (2012).
- [21] P. Adamson *et al.* (MINOS Collaboration), *Phys. Rev. Lett.* **107**, 021801 (2011).
- [22] R. Ospanov, Ph.D. thesis, University of Texas at Austin 2008.
- [23] C. Backhouse, D.Phil. thesis, Oxford University 2011.
- [24] S. Kopp, [arXiv:physics/0508001](https://arxiv.org/abs/physics/0508001).
- [25] K. Anderson *et al.*, Technical Report No. FERMILAB-DESIGN-1998-01, Fermilab (1998).
- [26] A zero field is expected due to symmetry around the beam axis. We measure a  $1 \times 10^{-2}$  mrad deflection of the primary proton beam, corresponding to  $B \cdot d\ell = 43$  Gm.
- [27] The window is defined by GPS time stamping each beam extraction and transmitting the time stamp over the internet to the FD, which buffers data to cope with latency.
- [28]  $\nu_\mu$  are assumed to oscillate with  $\Delta m^2 = 2.32 \times 10^{-3}$  eV<sup>2</sup> and  $\sin^2(2\theta) = 1$ .
- [29] The “ $p$  value,”  $p$ , is the probability, assuming our fit hypothesis is true, of obtaining data more incompatible with our hypothesis than the data we actually observe. See the discussion in Sec. 33.2.2 of [13].
- [30] S. Dytman, H. Gallagher, and M. Kordosky, [arXiv:hep-ex/08062119](https://arxiv.org/abs/hep-ex/08062119).
- [31] A. Bodek, H. S. Budd, and M. E. Christy, *Eur. Phys. J. C* **71**, 1726 (2011).
- [32] G. J. Feldman and R. D. Cousins, *Phys. Rev. D* **57**, 3873 (1998).